

# GAMMA-RAY EMISSION FROM TWO BLAZARS BEHIND THE GALACTIC PLANE: B2013+370 & B2023+336

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## ABSTRACT

B2013+370 and B2023+336 are two blazars at low-galactic latitude that were previously proposed to be the counterparts for the EGRET unidentified sources 3EG J2016+3657 and 3EG J2027+3429. Gamma-ray emission associated with the EGRET sources has been detected by the *Fermi* Gamma-ray Space Telescope, and the two sources, 1FGL J2015.7+3708 and 1FGL J2027.6+3335, have been classified as unidentified in the 1-year catalog. This analysis of the *Fermi*-LAT data collected during 31 months reveals that the 1FGL sources are spatially compatible with the blazars, and are significantly variable, supporting the hypothesis of extragalactic origin for the gamma-ray emission. The gamma-ray light curves are compared with 15 GHz radio light curves from the 40-m telescope at the Owens Valley Radio Observatory (OVRO). Simultaneous variability is seen in both bands for the two blazar candidates. The study is completed with the X-ray analysis of 1FGL J2015.7+3708 using *Swift* observations that were triggered in August 2010 by a *Fermi*-detected flare. The resulting spectral energy distribution shows a two-component structure typical of blazars. We also identify a second source in the field of view of 1FGL J2027.6+3335 with similar characteristics to the known LAT pulsars. This study gives solid evidence favoring blazar counterparts for these two unidentified EGRET and *Fermi* sources, supporting the hypothesis that a number of unidentified gamma-ray sources at low galactic latitudes are indeed of extragalactic origin.

*Subject headings:* Galaxies: individual (B2013+370, B2023+336) – Gamma rays: galaxies

## 1. INTRODUCTION

Blazars are a subclass of Active Galactic Nuclei (AGN) which host supermassive black holes in their centers and have relativistic jets pointing along the line of sight to the observer. They show broadband non-thermal emission extending from radio frequencies to gamma rays, with a typical two-component spectral energy distribution (SED, see, e.g. Sambruna et al. 1996). In the *Fermi* Large Area Telescope (*Fermi*-LAT) 1-year catalog (1FGL, Abdo et al. 2010b) 573 out of 1451 (39%) detected objects are associated with known blazars, making them the most numerous class of gamma-ray sources. Only a few bright blazars have shown gamma-ray variability correlated with other wavelengths, allowing a firm identification. Almost all blazar associations reported in the 1FGL catalog are based on positional coincidence of the *Fermi*-LAT gamma-ray source with the radio location of a preselected blazar. Most gamma-ray blazars

are detected at high Galactic latitudes, and only 14 are reported at  $|b| < 10^\circ$  (Abdo et al. 2010b). The same deficit of detections at low Galactic latitudes is reported for gamma-ray emitting AGNs, which includes blazars and also misaligned active galaxies. The first *Fermi*-LAT AGN catalog (1LAC, Abdo et al. 2010c) makes 51 associations with low-latitude AGNs from the VLBA calibrator survey, but a simple extrapolation assuming an isotropic distribution of gamma-ray emitting AGNs indicates that *Fermi*-LAT should be detecting  $\sim 150$  AGNs at  $|b| < 10^\circ$  (Abdo et al. 2010c).

This spatial anisotropy in the distribution of gamma-ray blazars and AGNs is mainly caused by a lack of extragalactic source catalogs covering low Galactic latitudes. Catalogs of candidate gamma-ray blazars are generally built by selecting compact flat-spectrum radio sources and identifying and classifying them as blazars through optical spectroscopy (see, e.g. Healey et al. 2008). Close to the Galactic plane, diffuse radio emission and confusion with local radio sources make candidate blazars difficult to select. Heavy optical extinction due to interstellar dust (Schlegel et al. 1998) further complicates AGN studies at low Galactic latitudes. The three pre-selected catalogs used to find blazar associations for 1FGL sources avoid the Galactic plane due to these observational difficulties (Healey et al. 2007, 2008; Massaro et al. 2009). Therefore, the deficit of gamma-ray blazars at low Galactic latitudes is mainly an artifact of the association procedure caused by our limited observational knowledge of blazars behind the Galactic plane. In fact, unbiased gamma-ray surveys like the *Fermi* 1-year catalog might be the best tools to identify the population of blazars at low Galactic latitudes that are hardly visible at other

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TABLE 1  
NON-1FGL SOURCES INCLUDED IN THE MODEL

Name	RA [deg]	DEC [deg]	$\sigma_{95\%}$ [deg]	TS	Fit	$\Gamma(E_c)$	Flux 1-100 GeV [ $10^{-9} \text{ cm}^{-2} \text{ s}^{-1}$ ]	2FGL association
J2020.0+4158	304.70	42.16	0.08	232	PL	$2.94 \pm 0.03$	$3.50 \pm 0.22$	J2020.0+4159
J2022.2+3840	305.55	38.68	0.13	117	PL	$3.13 \pm 0.04$	$2.45 \pm 0.16$	J2022.8+3843c <sup>a</sup>
J2017.4+3628	304.45	36.42	0.02	465	PL	$2.59 \pm 0.02$	$7.72 \pm 0.28$	J2018.0+3626
J2025.1+3342	306.26	33.70	0.05	289	PL	$2.94 \pm 0.03$	$3.25 \pm 0.14$	J2025.1+3341
J2028.3+3333	307.08	33.55	0.03	1112	PL Exp	$0.97 \pm 0.05$ (1370 $\pm$ 31 MeV)	$10.74 \pm 0.38$	J2028.3+3332

<sup>a</sup>associated with SNR G76.9+1.0 (Arzoumanian et al. 2011)

wavelengths.

The third EGRET Catalog (Hartman et al. 1999) listed 66 high-confidence associations with blazars, but none of them is at  $|b| < 10^\circ$ . Of the 80 reported sources at low Galactic latitudes, 74 were left unidentified. A re-evaluation of the correlation of flat-spectrum radio sources with northern EGRET sources suggested blazar identifications for 5 low-latitude unidentified sources (Sowards-Emmerd et al. 2003). Later detailed studies on individual EGRET unidentified sources suggested blazar counterparts for 3EG J2016+3657 and 3EG J2027+3429. In this paper we establish firm blazar associations for these two unidentified gamma-ray sources.

The blazar B2013+370 (RA =  $20^{\text{h}}15^{\text{m}}28.80^{\text{s}}$ , DEC =  $37^\circ10'58''$ ) is located at Galactic latitude  $b = 1.22^\circ$  in the Cygnus region. It was proposed as a counterpart for the EGRET unidentified source 3EG J2016+3657 in a comprehensive study of the X-ray sources in the field followed by optical spectroscopy of the candidate counterparts (Mukherjee et al. 2000; Halpern et al. 2001). B2013+370 was originally detected as a flat-spectrum radio source (Duin et al. 1975; Weiler & Shaver 1978). It appears compact down to milliarcsecond scales (Lee et al. 2008). There is no spectroscopic redshift, although the optical counterpart is manifest via its variability (Halpern et al. 2001). Massaro et al. (2009) classify B2013+370 as a blazar of uncertain type, and Véron-Cetty & Véron (2010) list it as a presumed BL Lac object. Very High Energy (VHE) gamma-ray observations by the Whipple 10-m telescope led to a 99% confidence level (c.l.) flux upper limit of  $F(E > 350 \text{ GeV}) < 2.0 \times 10^{-11} \text{ cm}^{-2} \text{ s}^{-1}$  assuming a photon index of 2.5 (Fegan et al. 2005).

B2023+336 (RA =  $20^{\text{h}}25^{\text{m}}10.84^{\text{s}}$ , DEC =  $33^\circ43'00''$ ), another blazar in the Cygnus region located at  $b = -2.37^\circ$ , was also proposed as counterpart for an EGRET unidentified source. Sowards-Emmerd et al. (2003) first suggested the association of 3EG J2027+3429 with B2023+336 after re-evaluating the correlation between EGRET unidentified sources and flat-spectrum radio sources. A later study supported the association by reporting significant flux variability from the blazar in the X-ray band (Sguera et al. 2004). Previously, other studies had suggested a young pulsar as counterpart for the EGRET source given its positional coincidence with an OB association (Romero et al. 1999; Zhang et al. 2000). At lower energies, B2023+336 has been detected as a compact flat-spectrum radio source (Griffith et al. 1990; Lee et al. 2008). Sowards-Emmerd et al. (2003) obtained a spectrum of the optical counterpart coincident with the radio position. The authors classified B2023+336 as a BL Lac object and derived a redshift of  $z = 0.219$  based

on H spectral lines. However, the signal-to-noise ratio of the optical spectrum is quite low due to heavy optical extinction, and the obtained redshift should be considered as tentative.

In the *Fermi* 1-year catalog, the blazars B2013+370 and B2023+336 are respectively found in the vicinity of the unidentified sources 1FGL J2015.7+3708 and 1FGL J2027.6+3335 (Abdo et al. 2010b).

In this paper we determine, discuss, and establish firm blazar associations for the gamma-ray sources 3EG J2016+3657 and 3EG J2027+3429 based on new gamma-ray, X-ray, optical and radio data.

## 2. GAMMA-RAY OBSERVATIONS

The LAT is the main instrument on-board the *Fermi* Gamma-ray Space Telescope. It is a multi-purpose telescope that surveys the gamma-ray sky between 100 MeV and 300 GeV, including the largely unexplored energy window between 10-100 GeV.

We selected photons from the first 31 months of the *Fermi* mission (Aug 2008 - Feb 2011). The analysis was performed using the *Fermi*-LAT *ScienceTools* software package (version v9r18). Diffuse class photons (Pass6 V3 IRF, Atwood et al. 2009) with  $E > 300 \text{ MeV}$  were selected in a  $17^\circ \times 17^\circ$  region around B2013+370, a Region of Interest (ROI) which also contains B2023+336.

We generated a model that includes all the sources in the *Fermi* 1-year catalog inside the ROI (46 1FGL sources total), and the Galactic and isotropic diffuse gamma-ray emission<sup>8</sup>. The 1FGL sources were modeled by a simple power-law spectrum of the form:  $dN/dE = N_0 (E/E_0)^{-\Gamma}$  where the photon index  $\Gamma$  and the normalization  $N_0$  were left free. The only exception is for the known gamma-ray pulsars in the ROI (Abdo et al. 2010d), which have been described by a power-law with exponential cutoff,  $dN/dE = N_0 (E/E_0)^{-\Gamma} e^{-E/E_c}$ , where the cutoff energy,  $E_c$ , was also left free. Known gamma-ray sources outside the ROI but within  $5^\circ$  from the edges were included in the model, but their spectral parameters were fixed to the 1FGL values. The isotropic diffuse background was modeled using the public spectral template (isotropic\_iem\_v02.txt) with a free normalization<sup>9</sup>. Galactic diffuse emission was also modeled by a template released by the *Fermi* Collaboration (gll\_iem\_v02.fit)<sup>8</sup>. Because the sources of interest lie close to the Galactic plane and the background is therefore

<sup>8</sup> This analysis was completed before the *Fermi* 2-year catalog was released, and only information from 1FGL was included in the analysis.

<sup>9</sup> <http://fermi.gsfc.nasa.gov/ssc/>

TABLE 2  
ANALYSIS RESULTS FOR THE GAMMA-RAY SOURCES STUDIED IN THIS WORK

Name	RA [deg]	DEC [deg]	$\sigma_{95\%}$ [deg]	TS	Fit	$\Gamma(E_c)$	Flux 1-100 GeV [ $10^{-9} \text{ cm}^{-2} \text{ s}^{-1}$ ]
1FGL J2015.7+3708	303.89	37.17	0.03	808	PL	$2.57 \pm 0.02$	$9.26 \pm 0.26$
J2025.1+3342	306.26	33.70	0.05	289	PL	$2.94 \pm 0.03$	$3.25 \pm 0.14$
J2028.3+3333	307.08	33.55	0.03	1112	PL Exp	$0.97 \pm 0.05$ ( $1.37 \pm 0.3 \text{ GeV}$ )	$10.74 \pm 0.38$

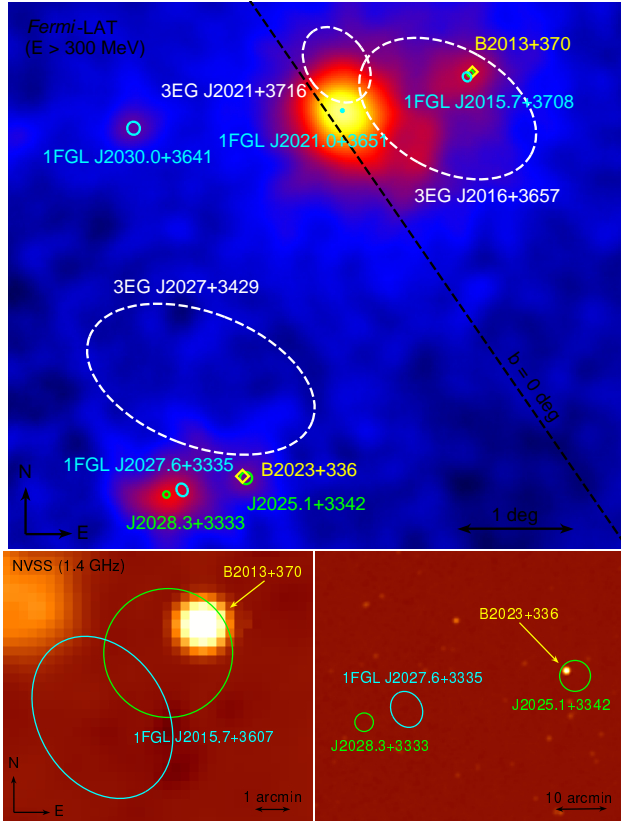


FIG. 1.— Top: *Fermi*-LAT counts map ( $E > 300 \text{ MeV}$ ) covering B2013+370 and B2023+336 where the galactic and diffuse gamma-ray emission have been subtracted. Bottom: Close-up NVSS counts maps (1.4 GHz, Condon et al. 1998) of B2013+370 (left) and B2023+336 (right). White dashed contours indicate 95% c.l. error ellipses for EGRET sources from Mattox et al. (2001). The error ellipses for the sources in *Fermi* 1-year catalog are shown in cyan. Green circles indicate the 95% c.l. error circles derived in this analysis, and yellow markers indicate the catalog positions of the two blazars.

quite heterogeneous, we weighted the Galactic diffuse emission by a power-law with free normalization and index, yielding a decrease of 11% on the normalization and a spectral hardening of 3% with respect to the public template. These additional degrees of freedom significantly improved the overall likelihood of the fit.

A binned likelihood analysis was performed using the `gtlike` tool. Using a residuals map, we identified additional sources not in the 1-year catalog, and included them in the model if they had  $TS > 25$ , equivalent to a significance  $\sigma \simeq 5$  (Mattox et al. 1996). The location of the additional sources was refined using `gtfindsrc`. Details on these additional sources are included in Table 1<sup>10</sup>.

To derive the spectra and light curves of the studied

sources the data were divided in time bins for the light curves and energy bins for the spectra, and the likelihood analysis was performed on each bin. To avoid having too many free parameters, the Galactic diffuse emission was fixed. The photon indices of all sources were also fixed, but their normalizations left free. The spectra were generated using the `likeSED` python module<sup>11</sup>. Spectral points in the spectra were required to have  $TS > 9$ . To keep the maximum number of data points in the light curves, flux points were derived when  $TS > 1$  as described in Abdo et al. (2009a). Otherwise 90% confidence level upper limits were calculated.

In our analysis we find that a single source at the location of 1FGL J2027.6+3335 does not represent the data well, and so it is removed from the model and replaced by two separate point-like sources (see Figure 1). A likelihood ratio test (LRT) confirms that the data favors the 2-source hypothesis. We compared a model with only the *Fermi* 1-year catalog source, 1FGL J2027.6+3335, to a model excluding the 1FGL source and including two sources instead: J2025.1+3342 and J2028.3+3333. The LRT favors the 2-source hypothesis with  $\chi^2/\text{dof} = 282/7$ , corresponding to  $\sim 16\sigma$ .

All results were cross-checked by an independent analysis, and are also found in agreement with a previous analysis of the same field of view including 27 months of data (Errando & Kara 2010).

In the following we describe more detailed results for these two individual gamma-ray sources, as well as 1FGL 2015.7+3708, all of which are also summarized in Table 2. All reported errors are statistical only. The estimated systematic uncertainty on the flux is 5% at 500 MeV, increasing to 20% at 10 GeV (Abdo et al. 2010b).

### 2.1. 1FGL J2015.7+3708

In this analysis 1FGL J2015.7+3708 is detected with  $TS = 808$ , equivalent to approximately  $28\sigma$ . The location of the gamma-ray source was recalculated using 31 months of data to be  $\text{RA} = 20^{\text{h}}15^{\text{m}}33.6^{\text{s}}$ ,  $\text{DEC} = 37^{\circ}10'12''$  with 95% error circle of  $1'8''$  that includes B2013+370. The time-averaged spectrum is well described with a power-law with photon index  $\Gamma = 2.57 \pm 0.02$ , and a measured integral flux above 1 GeV of  $(9.26 \pm 0.26) \times 10^{-9} \text{ cm}^{-2} \text{ s}^{-1}$ . Our result shows a softer spectrum than the one found in EGRET ( $\Gamma = 2.09 \pm 0.11$ , Hartman et al. 1999).

### 2.2. J2025.1+3342

In this analysis of 31 months of *Fermi*-LAT data we resolve a gamma-ray source, J2025.1+3342, towards the direction of 1FGL J2027.6+3335. It is located at  $\text{RA} = 20^{\text{h}}25^{\text{m}}02.4^{\text{s}}$ ,  $\text{DEC} = 33^{\circ}42'00''$  with 95% error circle

<sup>10</sup> All additional sources have been reported in the 2FGL catalog

<sup>11</sup> <http://fermi.gsfc.nasa.gov/ssc/data/analysis/user/>

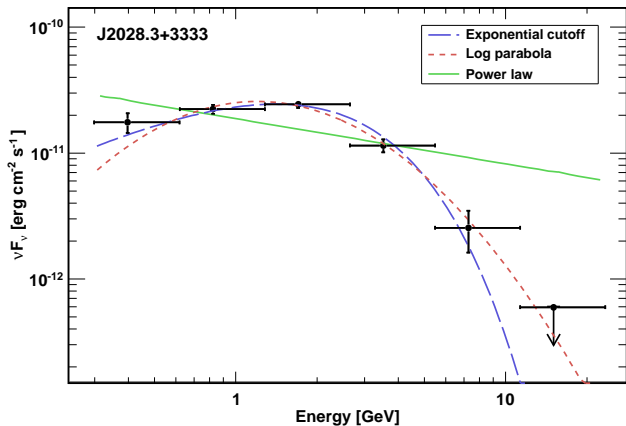


FIG. 2.— Gamma-ray spectrum of J2028.3+3333. The photon distribution The best-fit descriptions of the photon distribution in the whole energy range. The data was fit to a simple power-law, log-parabola and a power-law with exponential cut-off. The spectrum is best described by a power-law with exponential cutoff.

of 3' that includes the blazar B2023+336. J2025.1+3342 is detected with  $TS = 289$ , equivalent to approximately  $17\sigma$ . The energy spectrum is well described by a power-law with a photon index  $\Gamma = 2.94 \pm 0.03$  and integral flux above 1 GeV of  $(3.25 \pm 0.14) \times 10^{-9} \text{ cm}^{-2} \text{ s}^{-1}$ .

### 2.3. J2028.3+3333

In addition to J2025.1+3342 we resolve another, brighter gamma-ray emitter, J2028.3+3333, which is likely the primary contributor to 1FGL J2027.6+3335. The second, brighter source, J2028.3+3333, is detected with  $TS = 1112$ , or approximately  $33\sigma$ . As indicated in Errando & Kara (2010) the energy spectrum is best fit by a power-law with exponential cutoff where  $\Gamma = 0.97 \pm 0.05$ ,  $E_c = 1.37 \pm 0.03 \text{ GeV}$  and integral flux above 1 GeV of  $(10.74 \pm 0.38) \times 10^{-9} \text{ cm}^{-2} \text{ s}^{-1}$ . A LRT comparing a power-law, exponential cutoff and log-parabola spectral fit favors the exponential cutoff with a  $\chi^2/\text{dof} = 159/1$  ( $> 12\sigma$ ) with respect to a simple power-law (Figure 2). This source does not exhibit variability in gamma-rays. A  $\chi^2$  test on a 7-day binned light curve (not shown) gives a probability of the source being variable of  $P = 3\%$ .

### 2.4. Comparison with Fermi 2FGL catalog

After our gamma-ray analysis was completed, a first version of the *Fermi*-LAT 2-year catalog (2FGL) was released (Abdo et al. 2011c). The 2FGL catalog, which integrates over 24 months of LAT data, confirms the results obtained for 1FGL J2015.7+3708, J2025.1+3342, and J2028.3+3333, which are respectively associated with 2FGL J2015.6+3709, 2FGL J2025.1+3341, and 2FGL J2028.3+3332. The location and spectral index agrees within statistical errors for the three sources. Differences up to 20% in flux normalization are observed in 1FGL J2015.7+3708 and J2025.1+3342, which could be enhanced by the different integration time and the fact that both sources are variable. Another difference is that 2FGL lists a LogParabola function as best description for the energy spectrum of the three discussed sources, while our analysis preferred power-law spectra for the two blazars. This difference might be explained by the different energy cuts, as it is observed

TABLE 3  
*Swift*-XRT OBSERVATION LOG

Observation ID	Start time (UT) [yyyy-mm-dd hh:mm:ss]	XRT Exposure [sec]
00035639001	2006-07-12 05:29:01	2679.82
00035639002	2006-11-12 16:33:01	4551.96
00035639003	2006-11-17 04:09:01	7449.10
00035639004	2006-11-24 04:50:01	203.65
00041471001	2010-08-05 18:18:48	994.50
00041471002	2010-08-06 01:13:38	7424.83
00041471003	2010-08-22 05:35:00	4360.92
00041471004	2010-08-30 01:08:01	4216.07

that the biggest deviation from a power-law spectrum in the 2FGL sources happens to be in the 100-300 MeV bin, which was excluded from our analysis. The other reported discrepancies could be caused by multiple analysis differences: models for the Galactic and isotropic diffuse emission, time periods covered (31 months in this study vs. 24 in the 2FGL catalog), event calibrations (P6V3 vs. P7V6), instrument response functions, and energy cuts ( $E > 300 \text{ MeV}$  vs.  $E > 100 \text{ MeV}$ ), or versions of the ScienceTools. Finally, we note that 2FGL J2015.6+3709 and 2FGL J2025.1+3341 are respectively associated to the radio counterparts of B2013+370 and B2023+336, while 2FGL J2028.3+3332 is left unassociated.

### 3. X-RAY AND OPTICAL OBSERVATIONS

The *Swift* Gamma-Ray Explorer is designed to make prompt observations of Gamma-ray Bursts and afterglows, but also performs pointed observations. For this analysis, we use data from the *Swift* X-ray Telescope (XRT), which is sensitive from 0.2 to 10 keV (Gehrels et al. 2004).

During August 2010, four *Swift* pointed observations of B2013+370 (1FGL J2015.7+3708) were triggered by a *Fermi* detected flare on 2010 August 4<sup>12</sup>. Archival data from observations taken in 2006 were also analyzed. We performed two separate analysis on the 2010 and 2006 observations. Table 3 summarizes the *Swift*-XRT observations of B2013+370. No *Swift* data on B2023+336 are available.

The XRT data were processed with standard procedures in `xrtpipeline`, adopting the standard filtering and screening criteria, using FTOOLS in the `Heasoft` package (v6.10)<sup>13</sup>. For both the 2006 and 2010 sets, we summed the data from the four individual observations. Due to the low count rate of this weak source ( $6.0 \times 10^{-2} \text{ counts s}^{-1}$  in 2010, and  $2.3 \times 10^{-2} \text{ counts s}^{-1}$  in 2006), we consider only photon counting data, and further select XRT event grades 0-12 (Burrows et al. 2005). Pile-up correction was not required. From the 2010 and 2006 counts maps, we extracted the source events from a circular region with radius 20 pixels (1 pixel =  $2.36''$ ), and also extracted background events within a 55-pixel radius circle that was offset from the source.

Ancillary response files were generated with `xrtmkarf`, accounting for the different extraction regions, vignetting and PSF corrections. We used current response matrix file v011 available from HEASARC Calibration

<sup>12</sup> <http://fermisky.blogspot.com/>

<sup>13</sup> <http://heasarc.gsfc.nasa.gov/>

Database<sup>14</sup>. The spectra were rebinned in order to have at least 20 counts per energy bin, and the spectral fitting were completed using *XSPEC v12.6.0q*.

The X-ray spectra were fit by a power-law  $F(E) = KE^{-\Gamma}$ , absorbed by the Galactic column density,  $N_{\text{H}}^{\text{Gal}} \sim 1.7 \times 10^{22} \text{ cm}^{-2}$ , which was determined from the data. The spectrum from the August 2010 observations resulted in a photon index  $\Gamma = 1.77^{+0.22}_{-0.20}$ , and normalization at 1 keV of  $K = 1.46^{+0.57}_{-0.39} \times 10^{-3} \text{ keV}^{-1} \text{ cm}^{-2} \text{ s}^{-1}$ . The spectrum from the 2006 observations showed a similar photon index and higher flux ( $\Gamma = 1.23^{+0.55}_{-0.40}$ ,  $K = 3.26^{+0.41}_{-1.50} \times 10^{-3} \text{ keV}^{-1} \text{ cm}^{-2} \text{ s}^{-1}$ ).

Archival optical observations, imaged on 2004 July 10, were available for the position of B2023+336 from the 2.4m Hiltner Telescope of the MDM Observatory with a thinned, backside-illuminated SiTe CCD having a spatial scale of  $0''.275$  per  $24\mu$  pixel. Three 5 minute exposures in each of the *B*, *R*, and *I* filters were obtained in seeing of  $0''.8$ , and were combined. An astrometric solution for the images was derived from the UCAC2 (Zacharias et al. 2004). The optical counterpart of B2023+336 was clearly detected in each band, at position R.A. =  $20^{\text{h}}25^{\text{m}}10^{\text{s}}.85$ , decl. =  $+33^{\circ}43'00''.3$  (J2000.0), which is consistent with the radio position. Magnitudes were calibrated with Landolt (1992) standard stars. Including both statistical and systematic errors, we find  $B = 24.12 \pm 0.14$ ,  $R = 21.07 \pm 0.04$ , and  $I = 19.43 \pm 0.04$ .

Galactic extinction is a significant factor in the direction of B2023+336,  $(l, b) = (73^{\circ}.13, -2^{\circ}.37)$ . The Hydrogen column density ( $N_{\text{H}} = 6.27 \times 10^{21} \text{ Kalberla et al. 2005}$ ) was converted to optical extinction,  $A_{\text{V}} = 2.83$ , using the linear relation described in Güver & Oumlzel (2009), and scaled to the selected bandpasses using ratios given in Schlegel et al. (1998). The extinction-corrected optical fluxes for the *B*, *R* and *I* band can be found in Table 4.

#### 4. RADIO OBSERVATIONS

Contemporaneous 15 GHz radio light curves were obtained at the radio positions of the blazars B2013+370 and B2023+336 with the 40-m telescope at the Owens Valley Radio Observatory (OVRO). Since late 2007, about a year before the launch of *Fermi*, a large-scale, fast-cadence 15 GHz radio monitoring program has been ongoing using the 40-m telescope (Richards et al. 2011). The OVRO 40-m program began with the 1158 northern ( $\delta > -20^{\circ}$ ) sources from the Candidate Gamma-ray Blazar Survey (Healey et al. 2008), and now includes over 1500 sources including all the northern *Fermi* 1LAC blazar associations (Abdo et al. 2010c). The two sources in this analysis were added to the OVRO 40-m program in August, 2008. Program sources are each observed twice per week in total intensity with about 4 mJy (minimum) and 3% (typical) uncertainty. The absolute flux density scale is based on the Baars et al. (1977) value for 3C 286, 3.44 Jy at 15 GHz, with about 5% scale uncertainty which is not reflected in the error bars presented here.

Radio light curves of B2013+370 and B2023+336 are respectively shown in Figure 3 and 4, and the time-averaged radio fluxes are reported in Table 5 and 4.

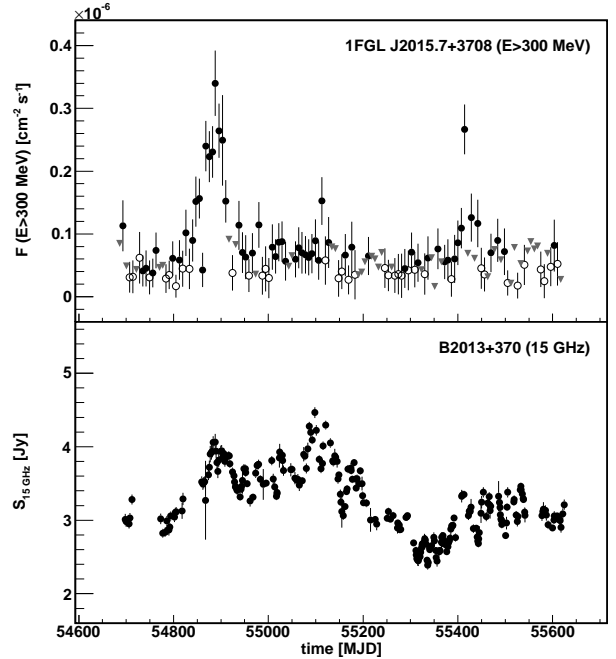


FIG. 3.— Top: Gamma-ray light curve at  $E > 300 \text{ MeV}$  in 7-day bins for 1FGL J2015.7+3708 (B2013+370). Open circles indicate fluxes for time periods where the source was detected with  $TS < 4$ , and grey triangles represent 90% c.l. upper limits ( $TS < 1$ ). Bottom: OVRO radio light curve at 15 GHz for same 31 month period.

#### 5. LOCATION OF THE GAMMA-RAY SOURCE

A detailed analysis of the *Fermi*-LAT data in the direction of B2013+370 and B2023+336 was presented in Section 2. The resolved gamma-ray sources and their locations with respect to the radio positions of the blazars are shown in Figure 1.

The possible association of the EGRET unidentified source 3EG J2016+3657 with the blazar B2013+370 was extensively discussed in Mukherjee et al. (2000) and Halpern et al. (2001). The authors examined X-ray images of the EGRET error box and conducted optical spectroscopy of all candidate counterparts, finding the blazar B2013+370 the only probable counterpart. With an 11-month exposure, the *Fermi*-LAT collaboration reported a gamma-ray excess (1FGL J2015.7+3708, Abdo et al. 2010b) associated with 3EG J2016+3657. The blazar B2013+370 lies just outside the 1FGL 95% error ellipse. A slight displacement of the *Fermi* reconstructed position could be caused by the high level Galactic diffuse emission in the region and the presence of a bright gamma-ray pulsar only  $1.2^{\circ}$  away (PSR J2021+3651, Abdo et al. 2010d). In the analysis of *Fermi*-LAT data presented here, we recalculated the location of the gamma-ray source with 3 times better photon statistics (see Table 2), and found the radio position of B2013+370 compatible with the reconstructed gamma-ray location (Figure 1). The newly available *Fermi* 2-year catalog confirms this association. Recently, the VERITAS collaboration reported the detection of a VHE gamma-ray source compatible with the nearby pulsar wind nebula CTB 87 at  $E > 650 \text{ GeV}$  (Aliu 2011). The emission is steady and significantly displaced from B2013+370, suggesting a different origin.

The blazar B2023+336 was proposed as the

<sup>14</sup> <http://heasarc.nasa.gov/docs/heasarc/caldb>

TABLE 4  
B2023+336 SPECTRAL POINTS

Instrument	Frequency [Hz]	$\nu F_\nu$ [ $10^{-12}$ erg s $^{-1}$ cm $^{-2}$ ]
OVRO	$1.5 \times 10^{10}$	$0.428 \pm 0.013$
MDM	$3.75 \times 10^{14}$	0.671
	$4.61 \times 10^{14}$	0.420
	$6.82 \times 10^{14}$	0.229
<i>Fermi</i> -LAT	$1.02 \times 10^{23}$	$21.2 \pm 2.5$
	$3.25 \times 10^{23}$	$6.19 \pm 1.13$
	$1.04 \times 10^{24}$	$3.31 \pm 0.79$
	$3.32 \times 10^{24}$	$< 1.01$
	$1.06 \times 10^{25}$	$< 1.29$

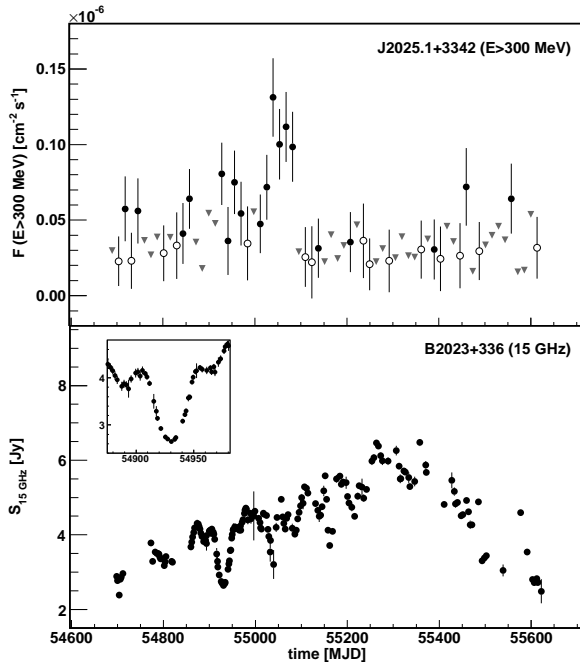


FIG. 4.— Top: Gamma-ray light curve at  $E > 300$  MeV in 14-day bins for J2025.1+3342 (B2023+336). Open circles indicate fluxes for time periods where the source was detected with  $TS < 4$ , and grey triangles represent 90% c.l. upper limits ( $TS < 1$ ). Bottom: OVRO radio light curve at 15 GHz for same 31 month period. The inset shows a zoom-in of the radio light curve around MJD 54930, when a flux modulation tentatively associated with an extreme scattering event is observed.

counterpart for the EGRET unidentified source 3EG J2027+3429 (Sowards-Emmerd et al. 2003; Sguera et al. 2004). The *Fermi*-LAT 11-month catalog reports an unidentified source (1FGL J2027.6+3335) outside the EGRET error ellipse. The position of the 1FGL source is  $32'$  away from the radio position of B2023+336. In our analysis presented in Section 2 we resolve a component of 1FGL J2027.6+3335: J2025.1+3342, spatially associated with the radio location of B2023+336 (Figure 1). The recent release of the *Fermi* 2-year catalog confirms this association.

## 6. FLUX VARIABILITY AND CORRELATIONS

Variability in gamma-ray sources allows an identification of the correct counterpart when correlated multi-wavelength variability is found. In addition, stand-alone gamma-ray variability can also support the affiliation of an unidentified source to an intrinsically variable source

class (see, e.g. Nolan et al. 2003). Flux variability is one of the main observational characteristics of blazars in the gamma-ray band (Abdo et al. 2010f), whereas only a handful of non-AGN sources have shown significant gamma-ray variability: 5 gamma-ray binaries (Abdo et al. 2010b; Corbet et al. 2011; Abdo et al. 2011b), V407 Cygni (Abdo et al. 2010e), and the Crab Nebula (Tavani et al. 2011; Abdo et al. 2011a).

The light curves of the gamma-ray sources associated with B2013+370 and B2023+336 exhibit some degree of variability. We estimated the gamma-ray variability with a  $\chi^2$  test on the light curve, including all flux points with  $TS > 1$  as in Abdo et al. (2009b). The 7-day binned light curve of 1FGL J2015.7+3708 (B2013+370) shown in Figure 3 exhibits clear flaring episodes, and significant evidence of the flux being variable ( $P > 99.9\%$ ). J2025.1+3342 (B2023+336) is a weaker gamma-ray emitter. We calculate a probability of its flux being variable of  $P = 99\%$  based on the 14-day binned light curve shown in Figure 4. This cannot be considered conclusive evidence of the source being variable. However, as noted in Abdo et al. (2010c), given the same fractional flux variation it is harder to find significant variability in faint sources like J2025.1+3342 than it is for bright sources because of the strong dependence of the variability test on the statistical flux uncertainties.

A possible correlation between the gamma-ray and radio light curves is studied by constructing a discrete correlation function (DCF, Edelson & Krolik 1988) shown in Figure 5. The statistical significance of the cross-correlation is investigated with Monte Carlo simulations. It is assumed that the light curves can be described as noise processes with a power-law power spectral density ( $\propto f^\beta$ ). The adopted power-law exponents are -1.5 for gamma-rays (Abdo et al. 2010f) and -2.0 for radio (Chatterjee et al. 2008). Using the recipe prescribed by Timmer & Koenig (1995), a pair of independent light curves is generated with a daily sample rate. These light curves are sampled at the same times as the actual radio and gamma-ray light curves and the values are perturbed by adding Gaussian white noise with zero mean and variance that matches the observational errors. The resulting mock light curves are then cross-correlated using the method described in Edelson & Krolik (1988). The process is repeated 10000 times, and the results provide an estimate of the distribution of random cross-correlations for each time lag.

The results of the cross-correlation tests are summarized in Figure 5, which shows the DCF for the data as black points and the 1- $\sigma$  (red), 2- $\sigma$  (yellow) and 3- $\sigma$  (green) contours for the distribution of random cross-correlations. Using these values we find no highly significant cross-correlation. The most prominent peaks are located at -10 days (radio lagging) with an 88.0% significance for B2013+370 and a peak at -240 days with a 96.7% significance for B2023+336. Additionally, the auto-correlation functions (not shown) were calculated separately for the gamma-ray and radio light curves, without finding any significant peak that could suggest periodic behavior.

Although a statistically significant correlation between the gamma-ray and radio light curves would provide a definite identification, such identifications based on correlated variability have only been established for very



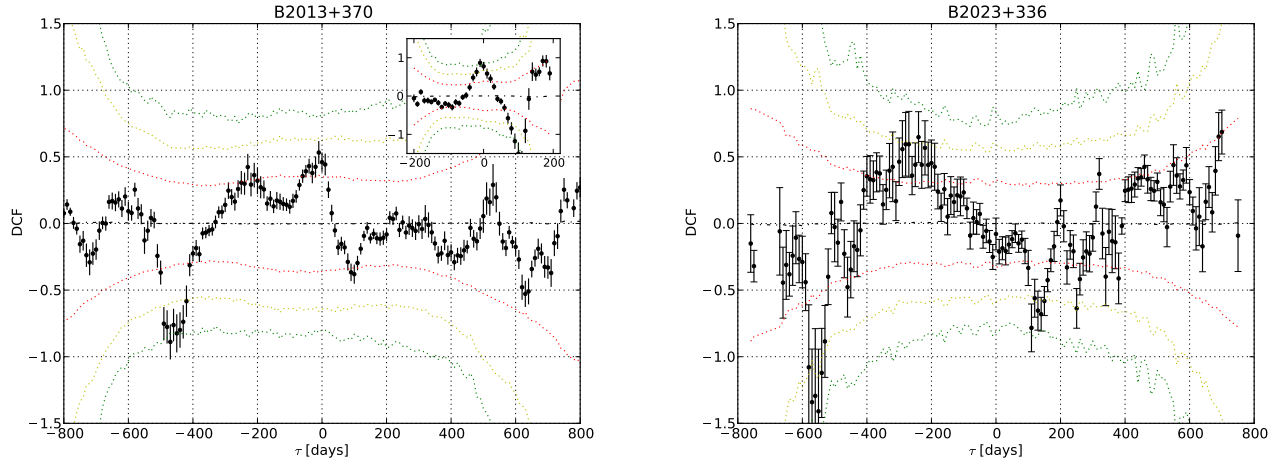


FIG. 5.— Discrete correlation function (DCF) between the gamma-ray and radio light curves for B2013+370 (left) and B2023+336 (right). Positive time lags indicate that radio activity precedes activity in gamma-rays. The significance of the correlations is illustrated by 1- $\sigma$  red, 2- $\sigma$  yellow and 3- $\sigma$  green dotted lines calculated from simulated uncorrelated data assuming a simple power-law power spectral density with exponent -2.0 for radio and -1.5 for gamma-ray light curves. The inset DCF for B2013+370 during the gamma-ray flare in MJD 54750-55000.

TABLE 5  
B2013+370 SPECTRAL POINTS

Instrument	Frequency [Hz]	Average		2010 August	
		$\nu F_\nu$ [ $10^{-12}$ erg s $^{-1}$ cm $^{-2}$ ]		$\nu F_\nu$ [ $10^{-12}$ erg s $^{-1}$ cm $^{-2}$ ]	
OVRO	$1.5 \times 10^{10}$	$0.485 \pm 0.01$		$1.5 \times 10^{10}$	$0.453 \pm 0.009$
<i>Swift</i> -XRT	$5.11 \times 10^{17}$	$1.43 \pm 0.10$		$5.84 \times 10^{17}$	$0.74 \pm 0.09$
	$8.42 \times 10^{17}$	$2.16 \pm 0.15$		$7.10 \times 10^{17}$	$2.15 \pm 0.14$
	$1.19 \times 10^{18}$	$3.05 \pm 0.22$		$9.66 \times 10^{17}$	$3.05 \pm 0.24$
	$1.51 \times 10^{18}$	$3.72 \pm 0.27$			
<i>Fermi</i> -LAT	$9.81 \times 10^{22}$	$36.1 \pm 10.3$		$1.05 \times 10^{23}$	$68.8 \pm 13.3$
	$2.25 \times 10^{23}$	$29.1 \pm 3.6$		$3.35 \times 10^{23}$	$47.7 \pm 10.4$
	$5.16 \times 10^{23}$	$15.8 \pm 2.1$		$1.07 \times 10^{24}$	$24.0 \pm 9.7$
	$1.18 \times 10^{24}$	$7.87 \pm 1.54$		$3.42 \times 10^{24}$	$< 10.9$
	$2.71 \times 10^{24}$	$3.64 \pm 1.29$		$1.09 \times 10^{25}$	$< 33.2$
	$6.22 \times 10^{24}$	$< 4.53$			
	$1.43 \times 10^{25}$	$< 8.35$			

bright gamma-ray blazars during short-lived flares (see, e.g. Wehrle et al. 1998; Iafrate et al. 2009; Abdo et al. 2010a). If we only include the prominent gamma-ray flare of B2013+370 in 2009 the value of the cross-correlation peak at -10 days lag clearly increases, as shown in Figure 5. Assessing the statistical significance for very short light curves is difficult because the noise properties used for the statistical test are derived from longer time series and might not be appropriate for short periods of time, especially when these have been selected because of unusual source activity, a prominent gamma-ray flare in this case. The fact that the significance decreases as more data are included indicates that no robust claims about the physical significance of the apparent correlations can be made using time series dominated by a single event.

Intrinsic long-term flux correlations might be washed out by the presence of extrinsic variability affecting only one of the frequency bands. Spangler & Cordes (1998) showed that observations of extragalactic sources in the line of sight to B2013+370 and near B2023+336 show heavy scattering due to interstellar plasma turbulence

in the direction of the Cygnus OB1 association. This mechanism would produce extrinsic variability in the radio band but not in gamma-rays. Disentangling intrinsic and extrinsic variability in the radio light curves is a difficult problem, but a particular feature on the light curve of B2023+336 around MJD 54930 strongly suggests the possibility of an extreme scattering event (Fiedler et al. 1987). Although extreme scattering events have only been detected at lower radio frequencies, the location of the source at low Galactic latitude and the characteristic symmetric flux decrease expected when an interstellar plasma lens moves through the line of sight, with the small flux increases at both ends of the event being caused by focusing of radio waves at the edges of the plasma cloud, strongly support this hypothesis (Clegg et al. 1998). To explore the effect of excluding the candidate extreme scattering event from the radio light curve of B2023+336, we use the same parameters as for the original analysis but exclude the radio data from MJD 54900 through MJD 54975. We find a slightly higher significance for the cross-correlation of 98.8% at the same time lag. This supports the suggestion that radio-specific

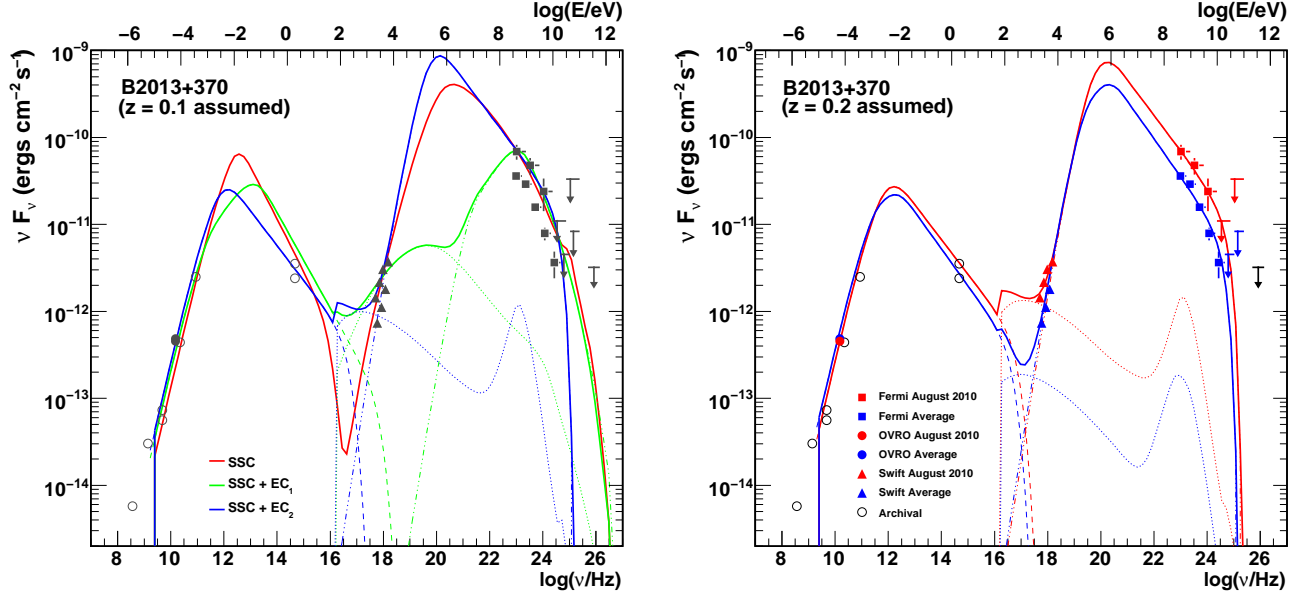


FIG. 6.— Spectral energy distribution for B2013+370. Left: Filled points (this work) correspond to OVRO (15 GHz), Swift XRT (2–10 keV) and *Fermi*-LAT data ( $E > 300$  MeV). Open circles show archival radio data obtained from NED,  $R$ -band magnitudes from Halpern et al. (2001), and an  $E > 350$  GeV upper limit from Fegan et al. (2005). Right: The blue (average) points consist of XRT data from 2006, and an average over the 31 months of Fermi and OVRO data. The red (August 2010) points consist of the Fermi and OVRO data from the entire month of August 2010, when there was a flare in gamma-rays, which triggered four XRT observations. Solid colored lines indicate the total predicted SED by the different tested models assuming a redshift of  $z = 0.1$  (left) and  $z = 0.2$  (right). Dashed lines show the Synchrotron emission, dotted lines the SSC component, and dot-dashed lines the EC.

TABLE 6  
SED MODELING PARAMETERS

Parameter	Symbol	SSC	SSC+EC <sub>1</sub>	SSC+EC <sub>2</sub>
Electron power	$L_e$ [erg s <sup>-1</sup> ]	$4.6 \times 10^{46}$	$4.7 \times 10^{44}$	$1.0 \times 10^{44}$
Electron low-energy cutoff	$\gamma_{min}$	$4.0 \times 10^3$	$1.5 \times 10^3$	$1.4 \times 10^2$
Electron high-energy cutoff	$\gamma_{max}$	$3.0 \times 10^5$	$3.0 \times 10^5$	$2.5 \times 10^4$
Electron injection index	$q_e$	3.2	3.1	2.8
Equipartition fraction	$L_B/L_e$	$1.58 \times 10^{-5}$	$6.49 \times 10^{-2}$	1.00
Blob radius	$R_b$ [cm]	$7.0 \times 10^{17}$	$6.0 \times 10^{16}$	$1.1 \times 10^{16}$
Magnetic field	$B$ [G]	0.001	0.1	1.0
Bulk Lorentz factor	$\Gamma$	20	15	15
Observing jet angle	$\theta_{obs}$ [deg]	2.87	3.82	3.82
Redshift (assumed)	$z$	0.1	0.1	0.1
External radiation energy density	$u_{ext}$ [erg cm <sup>-3</sup> ]	...	$2.0 \times 10^{-7}$	$2.5 \times 10^{-7}$
External radiation temperature	$T_{ext}$ [K]	...	$10^3$	$10^2$

extrinsic effects reduce the observable cross-correlation.

## 7. SPECTRAL ENERGY DISTRIBUTION

The SED for B2013+370 (Figure 6, data points in Table 5) shows a two component structure characteristic of blazars. The relatively hard X-ray spectral index ( $\alpha_X = 1.2 - 1.8$ ) and the dominance of the gamma-ray power output are typically observed in low-frequency-peaked BL Lacs (LBLs) or in flat spectrum radio quasars (FSRQs, Fossati et al. 1998).

We attempted to describe the multi-wavelength SED of B2013+370 with the one-zone leptonic emission model described in Böttcher & Chiang (2002) and Acciari et al. (2009). The model-predicted SEDs are shown in Figure 6, and the assumed model parameters are summarized in Table 6. The redshift of the source was assumed to be 0.1 throughout the modeling process. A pure synchrotron self-Compton model (SSC, e.g. Maraschi et al. 1992; Bloom & Marscher 1996) gives a poor description

of the SED and requires a very low magnetic field, resulting in a strongly particle dominated jet ( $L_B/L_e \sim 10^{-5}$ , with  $L_B$  and  $L_e$  being the magnetic and particle power). This is not surprising: SSC models typically fail to reproduce the high-energy component of Compton dominated blazars (LBLs, and specially FSRQs) unless the magnetic field is reduced to unphysically low values. This is addressed in external Compton models (EC, e.g. Dermer & Schlickeiser 1993; Sikora & Madejski 2000) by adding a population of low energy photons external to the jet. FSRQs are often described with leptonic models where SSC is dominating in the X-ray band and EC explains the gamma-ray output (e.g. Hartman et al. 2001). We explore this scenario by adding an external radiation field with temperature  $T_{ext} = 1000$  K. The resulting set of parameters (SSC+EC<sub>1</sub>) is closer to the typical values encountered in blazar SED modeling, but still needs low sub-equipartition magnetic fields and fails to reproduce the observed X-ray spectral slope. The best fit



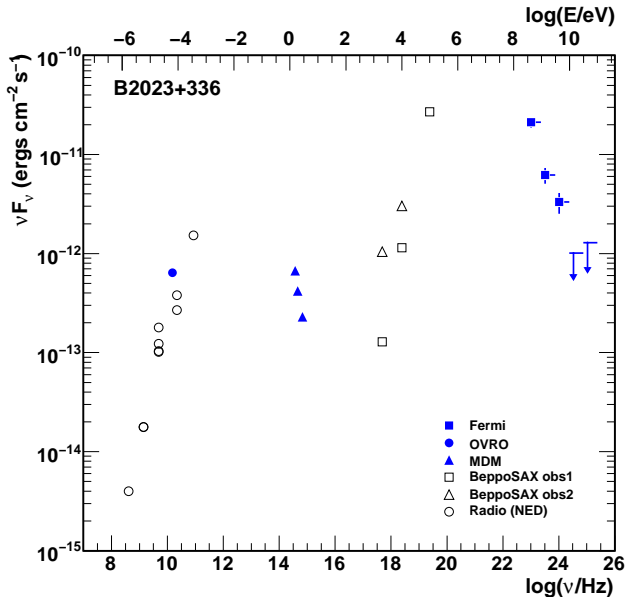


FIG. 7.— Spectral energy distribution for B2023+336 using OVRO (15 GHz), MDM (*B*, *R*, and *I* bands), and *Fermi* data ( $E > 300$  MeV). Open points indicate archival data from NED, and from two observations taken with BeppoSAX (1–100 keV) from Sguera et al. (2004).

(SSC+EC<sub>2</sub>) is achieved with both X-rays and gamma-rays being dominated by EC emission. Parameters in exact equipartition ( $L_B/L_e = 1$ ) can be achieved in this case, requiring a rather low temperature of the external radiation field ( $T_{ext} = 100$  K) that could originate from cold dust. We note that the relatively low frequencies of the photons produced in the first-order SSC ( $\sim 10^{17}$  Hz) make the second-order SSC (peaking at  $\sim 10^{23}$  Hz) become efficient. The lack of an  $E \gtrsim 50$  GeV detection, where the gamma-ray absorption by extragalactic background light is strongly redshift dependent, makes the assumed  $z = 0.1$  a non-critical model parameter. Similar best-fit solutions close to equipartition were found assuming  $z$  up to 1.5. A similar set of model parameters (assuming  $z = 0.2$  this time) can also describe the average and 2010 August states shown in Figure 6.

The SED for B2023+336, shown in Figure 7, shows the same blazar-like features discussed for B2013+370: two-component structure, hard X-ray spectrum, and dominant gamma-ray power. However, only archival non-contemporaneous optical and X-ray data are available. Given that B2023+336 is known to be variable at all frequencies, we do not attempt to model this source because of the lack of a simultaneous SED.

## 8. DISCUSSION AND CONCLUSIONS

Spatial association and the observed variability in the gamma-ray and radio bands allow us to establish a firm association between B2013+370 and the previously unidentified gamma-ray source 3EG J2016+3657 (1FGL J2015.7+3708), confirming the thesis of Mukherjee et al. (2000) and Halpern et al. (2001). A compiled SED, adding newly analyzed OVRO, *Swift*-XRT and *Fermi*-LAT data shows a two component structure that further supports the blazar association. The SED was successfully modeled with a leptonic one-zone model with both X-ray and gamma-ray power dominated by

external Comptonization of a low-temperature external radiation field. The gamma-ray dominated SED, hard X-ray spectrum, and preference for EC models point towards B2013+370 being an LBL or an FSRQ.

We also find strong evidence of gamma-ray emission from the blazar B2023+336, that was likely contributing to the unidentified gamma-ray source 3EG J2027+3429 (1FGL J2027.6+3335), as hypothesized by Sowards-Emmerd et al. (2003) and Sguera et al. (2004). Our analysis of the available *Fermi*-LAT data resolves two independent gamma-ray sources of very different character. J2025.1+3342 is spatially associated with the blazar B2023+336 and shows a hint of variability. The tentative observation of an extreme scattering event in the radio light curve of B2023+336 further supports the hypothesis of it being extragalactic in origin, the flux decrease being caused by a cloud of plasma in the galactic interstellar medium crossing the line of sight.

The other resolved component of 3EG J2027+3429 (J2028.3+3333) is characterized as a steady gamma-ray emitter with an exponentially cutoff energy spectrum, qualities which are typical of LAT detected pulsars. This result has recently been confirmed by Pletsch et al. (2011), who report the discovery of nine previously unknown gamma-ray pulsars using a blind search method. This discovery supports the hypothesis of Romero et al. (1999) and Zhang et al. (2000) that a pulsar could be responsible for the gamma-ray excess 3EG J2027+3429, or at least part of it.

Optical spectroscopy of the studied objects is complicated due to their low optical luminosity and the extreme dust absorption in the optical band characteristic of low Galactic latitudes. Future observations in large-aperture optical telescopes could establish a definitive spectroscopical identification of both blazars.

This work has been made possible for the first time because of the continuous, well sampled light curves in the gamma-ray and radio bands obtained over 2.5 years by *Fermi* and the fast-cadence OVRO blazar monitoring program. In completing this analysis, we establish firm associations for the blazars B2013+370 and B2023+336 with previously unidentified EGRET and *Fermi*-LAT sources. This confirms the gamma-ray band as likely the best tool to identify blazars located behind the Galactic plane, which are heavily absorbed at other wavelengths, and supports the hypothesis that a number unidentified gamma-ray sources at low Galactic latitudes are indeed of extragalactic origin.

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